

University of Groningen

Enhanced H₂O formation through dust grain chemistry in X-ray exposed environments

Meijerink, R.; Cazaux, S.; Spaans, M.

Published in:
Astronomy & astrophysics

DOI:
[10.1051/0004-6361/201117700](https://doi.org/10.1051/0004-6361/201117700)

IMPORTANT NOTE: You are advised to consult the publisher's version (publisher's PDF) if you wish to cite from it. Please check the document version below.

Document Version
Publisher's PDF, also known as Version of record

Publication date:
2012

[Link to publication in University of Groningen/UMCG research database](#)

Citation for published version (APA):

Meijerink, R., Cazaux, S., & Spaans, M. (2012). Enhanced H₂O formation through dust grain chemistry in X-ray exposed environments. *Astronomy & astrophysics*, 537, [A102]. <https://doi.org/10.1051/0004-6361/201117700>

Copyright

Other than for strictly personal use, it is not permitted to download or to forward/distribute the text or part of it without the consent of the author(s) and/or copyright holder(s), unless the work is under an open content license (like Creative Commons).

The publication may also be distributed here under the terms of Article 25fa of the Dutch Copyright Act, indicated by the "Taverne" license. More information can be found on the University of Groningen website: <https://www.rug.nl/library/open-access/self-archiving-pure/taverne-amendment>.

Take-down policy

If you believe that this document breaches copyright please contact us providing details, and we will remove access to the work immediately and investigate your claim.

Downloaded from the University of Groningen/UMCG research database (Pure): <http://www.rug.nl/research/portal>. For technical reasons the number of authors shown on this cover page is limited to 10 maximum.

Enhanced H₂O formation through dust grain chemistry in X-ray exposed environments

R. Meijerink^{1,2}, S. Cazaux², and M. Spaans²

¹ Leiden Observatory, Leiden University, PO Box 9513, 2300 RA Leiden, The Netherlands

² Kapteyn Astronomical Institute, PO Box 800, 9700 AV Groningen, The Netherlands
 e-mail: meijerink@astro.rug.nl

Received 14 July 2011 / Accepted 6 November 2011

ABSTRACT

Context. The ultraluminous infrared galaxy Mrk 231, which shows signs of both black hole accretion and star formation, exhibits very strong water rotational lines between $\lambda = 200\text{--}670\ \mu\text{m}$, comparable to the strength of the CO rotational lines. High-redshift quasars also show similar CO and H₂O line properties, while starburst galaxies, such as M 82, lack these very strong H₂O lines in the same wavelength range, but do show strong CO lines.

Aims. We explore the possibility of enhancing the gas phase H₂O abundance in X-ray exposed environments, using bare interstellar carbonaceous dust grains as a catalyst. Cloud-cloud collisions cause C and J shocks, and strip the grains of their ice layers. The internal UV field created by X-rays from the accreting black hole does not allow reforming the ice.

Methods. We determined the formation rates of both OH and H₂O on dust grains, with temperatures $T_{\text{dust}} = 10\text{--}60\ \text{K}$, using both Monte Carlo and rate equation method simulations. The acquired formation rates were added to our X-ray chemistry code, which allowed us to calculate the thermal and chemical structure of the interstellar medium near an active galactic nucleus.

Results. We derive analytic expressions for the formation of OH and H₂O on bare dust grains as a catalyst. Oxygen atoms arriving on the dust are released into the gas phase in the form of OH and H₂O. The efficiencies of this conversion due to the chemistry occurring on dust are near 30 percent for oxygen converted into OH and 60 percent for oxygen converted into H₂O between $T_{\text{dust}} = 15\text{--}40\ \text{K}$. At higher temperatures, the efficiencies decline rapidly. When the gas is mostly atomic, molecule formation on dust is dominant over the gas-phase route, which is then quenched by the low H₂ abundance. Here, it is possible to enhance the warm ($T > 200\ \text{K}$) water abundance by an order of magnitude in X-ray exposed environments. This helps explain the observed bright water lines in nearby and high-redshift ULIRGs and quasars.

Key words. astrochemistry – ISM: abundances – galaxies: active – dust, extinction – galaxies: ISM – galaxies: starburst

1. Introduction

Warm molecular gas ($T > 100\ \text{K}$) cools mainly through CO, H₂O, and H₂ ro-vibrational transitions (Neufeld & Kaufman 1993; Neufeld et al. 1995). The ground state lines of CO, up to $J = 6\text{--}5$, are extensively studied from the ground in many galaxies (see, e.g., Papadopoulos et al. 2010) up to a redshift $z \sim 0.04$. Unfortunately, the atmosphere makes it impossible to study CO transitions with $J > 8$ and H₂O rotational lines. It is these lines, however, that trace the warm molecular gas. Moreover, these lines are sensitive to extreme UV irradiation from star formation and X-rays due to supermassive black-hole accretion. Some high- J CO lines have been observed in bright galactic objects such as the Orion Bar and Sagittarius A (cf. Hollenbach & Tielens 1999), and the Class 0 object L1448-mm (Nisini et al. 1999). It is only recently that significant progress has been made in the study of the warm molecular interstellar medium (ISM) in nearby galaxies. Both high- J CO and rotational water lines have been observed abundantly after the launch of the *Herschel* Space Observatory, that operates between wavelengths $\lambda = 55\text{--}672\ \mu\text{m}$. The successful HerCULES open-time key program has observed 29 close (U)LIRGs with both SPIRE and PACS instruments, revealing

CO lines up CO $J = 14\text{--}13$ (while current follow-up PACS observations might detect even higher transition) and H₂O lines in Mrk 231 that have comparable strengths (van der Werf et al. 2010; González-Alfonso et al. 2010). These CO and H₂O lines are at least in part produced in regions of the galaxy exposed to X-rays, so-called XDRs (Maloney et al. 1996; Meijerink & Spaans 2005), but emission produced through shock excitation may also contribute, as observed for OH in Mrk 231 (Fischer et al. 2010). Moreover, *Herschel*-PACS also detected a lot of high- J CO lines, as well as H₂O and OH in embedded objects, such as DK Chamaeleontis (van Kempen et al. 2010) and in disk atmospheres, e.g., HD 100546 (Sturm et al. 2010), where X-rays also affect the chemical-thermal structure.

Understanding the excitation and chemistry of these important molecules is therefore key to studying the dominant physical processes in the ISM and star formation. CO is a linear molecule, and the radiation transfer models allow calculating far-infrared to millimeter spectra relatively easily for a given physical structure. Models predicting CO emission, where both the thermal-chemical structure and radiation transfer are calculated simultaneously, have more problems reaching agreement. The case for H₂O is even worse, thanks to its complex nature. High critical

densities $n_{\text{crit}} > 10^8\text{--}10^{10}\text{ cm}^{-3}$ in combination with high opacities, infrared-pumping, and maser activity, make an excitation calculation extremely challenging. Also, the excitation state in which water enters the gas phase after desorption or evaporation from a dust grain is very uncertain. A first guess would be to divide the excess energy over the levels using equipartition, hence 1/3 translational, 1/3 rotational, and 1/3 vibrational excitation, with a Boltzmann distribution for the sub-levels.

The formation of molecular hydrogen, H_2 , in the ISM has been a longstanding problem. At extremely low metallicities, when no dust is present, H_2 forms either through the H^- route ($\text{H}^- + \text{H} \rightarrow \text{H}_2 + \text{e}^-$; Glover 2003; Launay et al. 1991) or through a three-body reaction ($\text{H} + \text{H} + \text{H} \rightarrow \text{H}_2 + \text{H}$). The three-body reaction dominates at very high densities ($n > 10^9\text{ cm}^{-3}$; Galli & Palla 1998). For normal ISM conditions, H_2 is predominantly formed on dust grains, and much progress has been made over the past two decades, both experimentally and theoretically, in understanding this (much faster) formation of H_2 on the surfaces of grains (Pirronello et al. 1999; Katz et al. 1999; Cazaux & Tielens 2002; Zecho et al. 2002; Cazaux & Tielens 2004; Sha et al. 2005; Hornekar et al. 2006; Morisset et al. 2004; Cuppen et al. 2010b, and references therein). In general, the understanding is that CO is formed in the gas phase, but the origin of gas-phase water is not so clear-cut. It is possible to form water in the gas phase through either a chain of neutral-neutral reactions containing a number of temperature barriers or through ion-molecule reactions when the gas is moderately ionized ($x_e \sim 10^{-5}$).

The formation of water on interstellar dust can also be an efficient route in diffuse and dense clouds (Cuppen & Herbst 2007; Cuppen et al. 2010a). Several routes for forming water can be considered, involving successive hydrogenation of oxygen with either H_2 or H . Because species can stay on the dust, they can repetitively collide with each other and react, even if there is a strong barrier against the reaction. In this case, dust grains provide a favorable place for improbable reactions (with high barriers) to occur (Cazaux et al. 2010). Also, other routes for forming water, involving O_2 (Miyachi et al. 2008; Ioppolo et al. 2010) and O_3 (Mokrane et al. 2009) have been investigated and can be important at high dust temperatures ($T_{\text{dust}} > 20\text{--}30\text{ K}$). The molecules present on the surface can be released in the gas phase through several mechanisms: desorption upon formation if the reaction is exothermic; photo-dissociation of surface species leading to the desorption of products in the gas (i.e., photodissociation of H_2O leads to OH and H in the gas phase, Andersson et al. 2006); evaporation of the surface when the dust temperature is high enough (depends on the binding energy of the species). If many water molecules are formed and stay on the dust, then ices can be made. The bonds that hold the water molecules together are very strong (creating water clusters). As a result, when one or more monolayers of ice are formed on the grain, shocks will be the only effective way to remove the ice from grains (e.g., Draine 1995). It is thus only possible to have efficient formation of gas-phase water with grains as a catalyst, when the grains are (close to) bare.

Objects such as Mrk 231 have very violent environments in the central regions around the super-massive black hole. This particular object has a derived $8\text{--}1000\text{ }\mu\text{m}$ infrared luminosity $L_{\text{IR}} = 4.0 \times 10^{12} L_{\odot}$ (Boksenberg et al. 1977). A highly absorbed power-law, X-ray spectrum was observed by Braitto et al. (2004). It is in these extreme environments that we expect cloud-cloud collisions to cause C and J shocks and to strip the grains of their ice layers. The internal UV field created by X-rays from the accreting black hole does not allow any ice to be reformed.

The abundances of H_2O ($x_{\text{H}_2\text{O}} \sim 10^{-6}$, González-Alfonso et al. 2010) in the warm molecular gas of Mrk 231 are high. Here we study the effect of OH and H_2O formation on bare dust grains in environments with strong X-ray radiation fields, in order to investigate whether this process would be able to significantly boost the total production rate of these species in these environments. Our work differs from Hollenbach et al. (2009), since we consider an additional process for releasing molecular species into the gas phase, and also since our study only considers bare interstellar dust grains. Even though this paper focuses on an application to ULIRGs, it will have applications to other objects as well, such as X-ray irradiated disks around young stellar objects. The manuscript is organized in the following way. We first derive analytical expressions for the formation rates of OH and H_2O on dust grains. Then we compare these rates to the neutral-neutral, gas-phase formation processes. Finally, the effects of the grain surface chemistry on the XDR thermo-chemical structure are shown. Their implications for explaining recent observations with the *Herschel* Space Observatory are also discussed.

2. H_2O and OH formation rates

As mentioned earlier, gas-phase OH and H_2O can occur through neutral-neutral reactions or when the medium is moderately ionized, $x_e \gtrsim 10^{-5}$, through ion-molecule reactions. The neutral-neutral reactions $\text{OH} + \text{H}_2 \rightarrow \text{H}_2\text{O} + \text{H}$ and $\text{O} + \text{H}_2 \rightarrow \text{OH} + \text{H}$ are slow at low temperatures, because they have activation barriers (Wagner & Graff 1987) and only contribute at temperatures $T > 250\text{--}300\text{ K}$. The other route is $\text{H}^+ + \text{O} \rightarrow \text{O}^+ + \text{H}$, $\text{O}^+ + \text{H}_2 \rightarrow \text{OH}^+ + \text{H}$, $\text{OH}^+ + \text{H}_2 \rightarrow \text{H}_2\text{O}^+ + \text{H}$, $\text{H}_2\text{O}^+ + \text{H}_2 \rightarrow \text{H}_3\text{O}^+ + \text{H}$, followed by recombination $\text{H}_3\text{O}^+ + \text{e}^- \rightarrow \text{H}_2\text{O} + \text{H}$.

In our current X-ray chemistry code, only the grain surface reaction is taken into account for the formation of H_2 . To study the effect of OH and H_2O molecule formation on grain surfaces, similar reactions for these species will be added to the chemical network

$$\frac{dn(\text{OH})}{dt} = n(\text{O})v_{\text{O}}n_{\text{dust}}\sigma\epsilon_{\text{OH}}, \quad (1)$$

$$\frac{dn(\text{H}_2\text{O})}{dt} = n(\text{O})v_{\text{O}}n_{\text{dust}}\sigma\epsilon_{\text{H}_2\text{O}}, \quad (2)$$

where ϵ_{OH} and $\epsilon_{\text{H}_2\text{O}}$ are efficiencies for the formation of OH and H_2O , respectively, and v_{O} the velocity of oxygen atoms. The expression for the efficiencies is determined in the next section, where we calculate what fraction of an accreted oxygen atom is converted in gas-phase OH or H_2O . Two different grain size distributions are considered, an MRN (Mathis et al. 1977) distribution with $n_{\text{dust}}\sigma/n(\text{H}) = 10^{-21}\text{ cm}^{-2}$ and a Weingartner & Draine (2001) distribution with $2.8 \times 10^{-21}\text{ cm}^{-2}$. The MRN distribution considers a minimum dust grain size of $50\text{ }\text{\AA}$, while Weingartner and Draine (W&D) consider much smaller grains, which are actually PAHs and can be as small as $5\text{ }\text{\AA}$. The W&D grain size distribution, however, seems more appropriate to environments such as the central regions of Mrk 231, where dust grains can break into smaller particles due to, e.g., shocks. A larger cross section yields more enrichment of the medium with water and OH . The efficiencies are derived in the next section.

3. Grain surface chemistry

In this study, we consider carbon grains (graphite and carbonaceous). To describe the population of a certain species on a grain,

Table 1. Binding energies (in K) for carbon grains.

Physisorption		
H	550	Bergeron et al. (2008)
O	1390	Bergeron et al. (2008)
H ₂	600	Akai & Saito (2003)
OH	1360	Cuppen & Herbst (2007)
O ₂	1440	Cuppen & Herbst (2007)
H ₂ O	2000	Cuppen & Herbst (2007)
O ₃	2240	Lee et al. (2009)
HO ₂	2160	Cuppen & Herbst (2007)
H ₂ O ₂	2240	Cuppen & Herbst (2007)
Chemisorption:		
H	8500	Jelea et al. (2008)

we need to consider additional processes apart from the reactions that normally occur in the gas phase, such as neutral-neutral reactions. These are accretion to, evaporation from, tunneling and thermal hopping of atoms, and molecules on the surface of dust grains.

Species arriving from the gas phase arrive at a random time and location on the dust surface. The arrival time depends on the accretion rate, which can be written as

$$R_{\text{acc}} = n_X v_X \sigma S_X \quad (3)$$

where n_X and v_X are the densities and velocities, respectively, of the species X , and σ is the cross-section of the dust particle, and S_X is the sticking coefficient of the species with the dust. We consider $S_X = 1$. The species X can go back to the gas phase through evaporation. This rate can be written as

$$R_{\text{evap}} = \nu_i \exp(-E_X/kT) \quad (4)$$

with ν_i the oscillator factor of atom in site i (which is on the order of 10^{12} s^{-1} for physisorbed atoms), and E_X the binding energy of species X (see Table 1 for the adopted values). The binding energy E_X can be either weak (for atoms and molecules) or strong (only for atoms). In this study we consider binding energies similar to those in Cazaux et al. (2010). There are two types of interactions between the atoms and the surface: a weak one, called physisorption (van der Waals interaction), and a strong one, called chemisorption (covalent bond). Atoms from the gas phase can easily access the physisorbed sites and become physisorbed atoms. These weakly bound atoms can travel on the surface at very low dust temperatures, meeting each other to form molecules. Species are moved on dust grains via through tunneling or thermal hopping (Cazaux & Tielens 2004). The mobility of a species X moving from physisorbed sites to physisorbed sites can be written as

$$\alpha_{pp}(X) = \nu P_{pp}(X), \quad (5)$$

where $P_{pp}(X)$ is the probability that an atom moves from one physisorbed to another physisorbed site. In this study, we consider a range of surface temperatures $T > 15 \text{ K}$. At these temperatures, thermal hopping dominates and the probability that a species X goes from a physisorbed site to another physisorbed site can be written as

$$P_{pp}(X) = \exp(-E_{pp}(X)/kT), \quad (6)$$

where $E_{pp}(X)$ is the barrier between two physisorbed sites, which we assume to be two thirds of the binding energy E_X (see Table 1). Hydrogen atoms that are physisorbed can also enter chemisorbed sites through tunneling effects or thermal

hopping. Because of the high barrier against chemisorption Sha (2002), H atoms usually tunnel through the barrier to populate a chemisorbed site. The mobility for the H atoms can be written as

$$\alpha_{pc} = \nu P_{pc}, \quad (7)$$

where $P_{pc}(H)$ is the probability of a physisorbed H atom entering a chemisorbed site (Cazaux & Tielens 2004). Once two species meet at the same site, they can form a new species if the activation barrier for the formation can be overcome. According to the exothermicity of the reaction, the newly formed species is *directly released into the gas phase*. This probability, called f_{des} in this work, has been discussed in Cazaux et al. (2010), and is the most important route for releasing products to the gas phase. Species present on the surface can also receive a UV photon and become photo-dissociated. In X-ray exposed environments, an internal UV field is created by excitation of Lyman-Werner bands, and about 30 percent of the locally absorbed X-ray is converted into UV (Maloney et al. 1996). The photodissociation of H₂O on the dust can lead to the release in the gas phase of the two products of the dissociation, OH and H. This mechanism has a very small chance of occurring (2 percent, Andersson & van Dishoeck 2008). Although this process is much faster than thermal desorption (the grains are too cold to make thermal desorption efficient), it is not the main contributor in delivering species to the gas phase. On grains where ice layers are present, the desorption upon formation is less prominent, because the species are more strongly bound to the surface. In those environments, photodesorption is the dominant route for releasing species into the gas phase. More details about the adopted reaction rates can be found in Cazaux et al. (2010). Since we are interested in the oxygen chemistry, we consider H, H₂, O, OH, O₂, H₂O, O₃, HO₂, and H₂O₂ as grain surface species in the chemical network.

Two different numerical methods are performed to follow the formation of species on dust grain surfaces: (1) a Monte Carlo (MC) and (2) a rate equation method. The first is performed at several fixed dust temperatures. These calculations are very time consuming, and are used to check for which conditions the chemistry occurring on the dust enters the stochastic regime. The rate equation method is used to simulate the chemistry occurring on dust for a wide range of dust and gas temperatures, gas phase chemical composition, and radiation fields. This method is orders of magnitude faster than MC simulations.

3.1. Monte Carlo method

We use step-by-step MC simulations to follow the chemistry occurring on dust grains. The dust grains are divided into square lattices of 10×10 adsorption sites. Each site on the grid corresponds to a physisorbed site and a chemisorbed site, so that the grain can be seen as two superimposed grids. Species originating in the gas phase arrive at a random time and location on the dust surface. This arrival time depends on the rate at which gas species collide with the grain (Eq. (3)). When a species arrives at a point on the grid, it can become physisorbed, or, if its chemisorption states exists and its energy is high enough to cross the barrier against chemisorption, it becomes chemisorbed. In our model, however, because of the high barrier to accessing chemisorbed sites, H atoms mostly arrive from the gas phase in physisorbed sites. The species present on the surface may return to the gas phase if they evaporate (Eq. (4)). The species that arrive at a location on the surface can move randomly

Table 2. Models parameters.

Env.	n_{H}	$x(\text{HI})$	$x(\text{H}_2)$	$x(\text{OI})$	G_0
1	10^3	1	10^{-3}	3.4×10^{-4}	$10^0, 10^1, 10^2, 10^3$
2	"	10^{-1}	0.45	"	"
3	"	10^{-2}	0.5	"	"
4	$10^{5.5}$	1	10^{-3}	"	$10^3, 10^4, 10^5, 10^6$
5	"	10^{-1}	0.45	"	"
6	"	10^{-2}	0.5	"	"

across the surface by means of tunneling effects and thermal hopping (Eq. (5)). Species present on the surface can also receive a UV photon and become dissociated at the rates shown in Table A.1. Once a chemical species present on the dust moves, adsorbs a photon, or meets another species, the next event that occurs to this species is determined along with its time of occurrence. Therefore, the events that concern every species on the dust are ordered by time of occurrence, and for each event that occurs, the next event for the concerned species is determined. In Fig. 1, we follow the fraction of oxygen that is released in the gas in the form of OH and H₂O for environments 1, 2, 4, and 5 (see Table 2 for adopted parameters). The parameters of these environments represent typical ambient low and high densities and radiation fields in the ISM of active galaxy centers, and are similar to those in Meijerink & Spaans (2005). For each environment, 50–60% (30–40%) of the oxygen arriving on the dust goes back into the gas phase in the form of H₂O (OH). At higher dust temperatures (~35–40 K), on the other hand, species coming on the dust evaporate very fast. Therefore, the efficiencies decrease dramatically since the grain surfaces are only sporadically covered by species.

3.2. Rate equation method

The rate equation method can be applied when there is a minimum coverage of 1 atom on the dust. This condition is satisfied if the grains are not too warm, that species can encounter each other before evaporating. By comparing our rate equation results with Monte Carlo simulations, we will define a range of dust temperatures for which rate equations are valid.

The rate equation method couples the different equations for the different species on the dust. The processes described in the previous subsection are dependent on the dust and gas temperature, UV field, and exothermicity of the reaction occurring on the dust. The equation for species X can thus be written as follows:

$$\begin{aligned} \frac{dn(X)}{dt} = & n_{\text{gas}}(X)v_X N_s - n(X)n(Y)\alpha_{pp}(X) \\ & + (1 - f_{\text{des}})n(Y)n(Z)\alpha_{pp}(Y) - R_{\text{evap}_X}n(X) \\ & - R_{\text{phot}_X}n(X) + R_{\text{phot}_Y}n(Y) \end{aligned} \quad (8)$$

where $\alpha_{pp}(X)$ and $\alpha_{pp}(Y)$ are the mobilities of the species X and Y (see Eq. (5)). The different terms of this equation present (1) accretion of species X from the gas phase; (2) the formation of the new species that involve species X ; (3) creation of the species X by the encountering of species Y and Z that stay on the surface; (4) evaporation of species X ; (5) photodissociation of X on the dust; and (6) the creation of X by photodissociation of species Y .

3.3. Efficiencies derived from the rate equation method

The formation efficiency is defined as the fraction of oxygen $f(\text{O})$ that is released in the gas phase in the form of another species. We studied this in a few different environments, chosen such that they resemble the conditions found in the XDR models of Meijerink & Spaans (2005). The gas temperature is fixed at a constant temperature $T = 100$ K, and dust temperatures ranging from $T_{\text{dust}} = 10$ –60 K. Four different radiation fields were considered at two fixed densities, $n = 10^3$ and $10^{5.5} \text{ cm}^{-3}$, in both the atomic and molecular regimes. The parameters used are summarized in Table 2.

In Figs. 2 and 3, we show the obtained efficiencies for OH and H₂O. The figures show that the efficiencies for OH and H₂O are roughly constant between dust temperatures $T_{\text{dust}} \sim 15$ –35 K. In this regime most of the oxygen goes to H₂O. For dust temperatures $T_{\text{dust}} \geq 35$ –45 K, the total efficiency for forming molecules is rapidly decreasing, and most of the atomic oxygen will not react to another species before it evaporates from the grain. Therefore, the dust surface chemistry can only be effective when the dust temperatures are not too high. We recall that the way to form OH and H₂O is through $\text{O} + \text{H} \rightarrow \text{OH}$ and $\text{OH} + \text{H} \rightarrow \text{H}_2\text{O}$. For slightly higher dust temperatures, the chance of an OH molecule evaporating from the grain becomes higher than for meeting another H atom. Therefore at temperatures $T_{\text{dust}} \geq 35$ –40 K, most of the oxygen atoms are incorporated into OH, and then ϵ_{OH} is highest.

Increasing the radiation field, has a similar effect as increasing the dust temperature, i.e., it results in a lower $\epsilon_{\text{H}_2\text{O}}/\epsilon_{\text{OH}}$ ratio, which is caused by photodissociation of H₂O on the surface. The strongest effect is seen in environments 3 and 6, where the dominant hydrogen fraction is contained in H₂. In this environment, the hydrogen accretion rates are lowest, and thus also have the lowest atomic surface density on grains. When most of the gas is in atomic form, the hydrogen accretion is such that at high dust temperatures most of the oxygen is converted into OH. However, when the gas become more and more molecular, the atomic hydrogen accretion rate onto the dust decreases, and a larger fraction of atomic oxygen ends up in O₂. In environment 3, this is even the most likely product at temperature $T_{\text{dust}} > 35$ K.

In Figs. 4 and 5, the efficiencies for the formation of O₂, O₃, OH₂, and H₂O₂ are shown. The efficiencies for these species show a peak at dust temperatures $T_{\text{dust}} = 35$ –40 K. The maximum efficiency for producing O₂ ranges from $\sim 10^{-2}$ in the atomic gas to $\sim 10^{-1}$ in molecular gas, and generally the efficiencies for other species are lower. At temperatures $T > 40$ K, O₂ efficiencies dominate those for OH, but this is the regime where the formation of molecular species is not efficient.

3.4. Validity of the rate equation method

To estimate the uncertainties in the MC calculation and to check the consistency between the MC and rate equation methods, we did two additional MC simulations at dust temperatures of 20, 30, and 35 (low density) and 40 (high density). After a time t of ~ 1000 –2000 years (low density) and ~ 5 –10 years (high density), the efficiencies are converged and oscillate around a stable value. At $T_{\text{dust}} = 20$ and 30 K, the H₂O efficiencies vary between 0.5 to 0.6, and those for OH between 0.3 to 0.4, so are very similar to the rate equation method. At the higher dust temperatures (35 and 40 K), the values obtained with the MC method are lower. The rate equation method overestimates the efficiencies, and we are entering the stochastic regime.

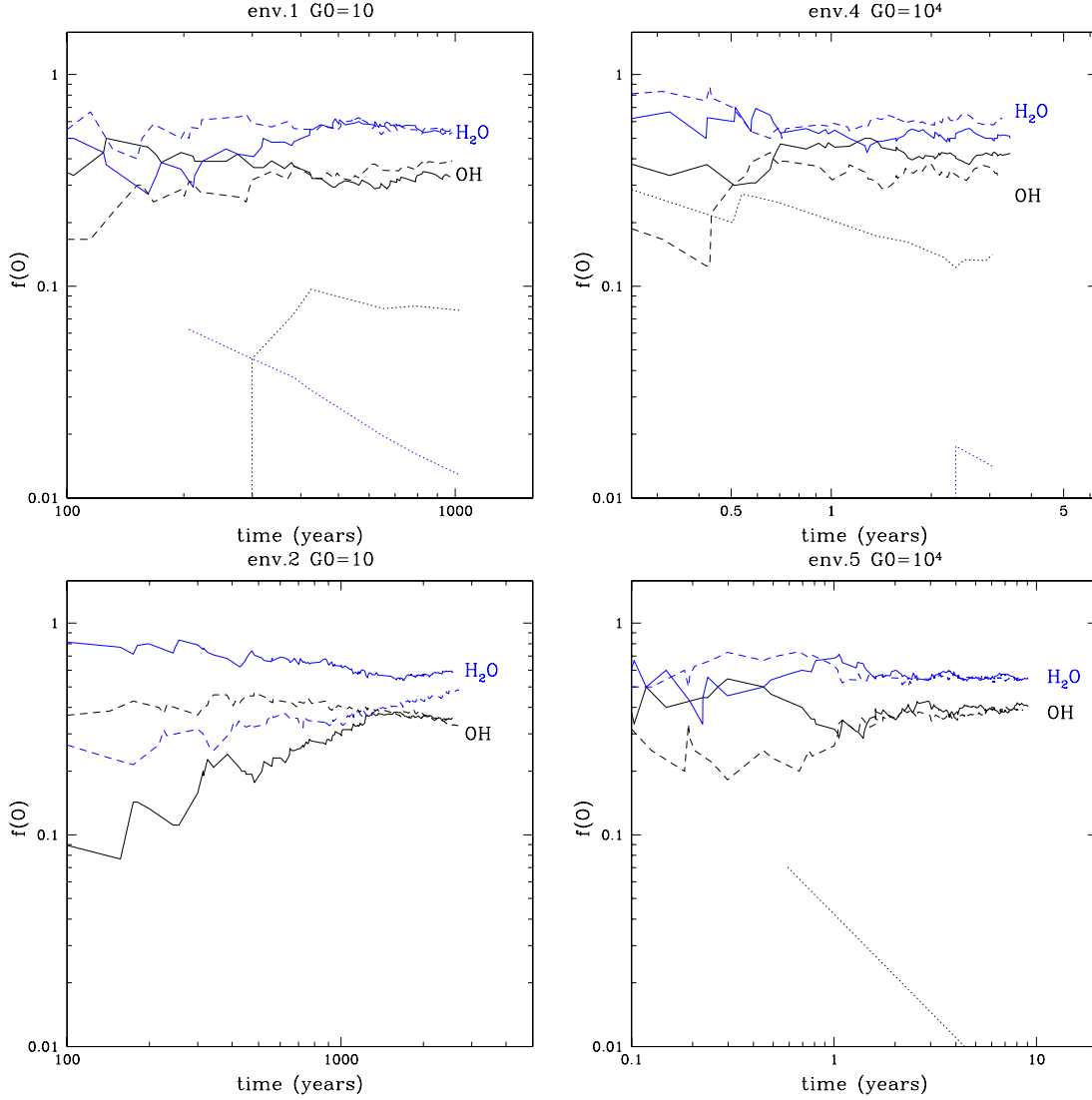


Fig. 1. Results of our Monte Carlo simulations on a 10×10 sites grain (30 \AA) for environments 1, 2, 4 and 5. The solid lines represent our results for $T_{\text{dust}} = 20 \text{ K}$, dashed lines for $T_{\text{dust}} = 30 \text{ K}$ and dotted lines $T_{\text{dust}} = 35 \text{ K}$ (left panel) 40 K (right panel).

We only concentrate on dust temperature $T_{\text{dust}} > 10\text{--}15 \text{ K}$. At very low dust temperatures, OH and H₂O are still able to form at a lower efficiency, but the chemistry is different then, for a number of reasons: (1) ices can form on the surface; (2) the chemistry is different because the binding energies are different; and (3) the surface of the grain will be saturated with H₂.

4. Analytical expressions for OH and H₂O formation efficiencies

The rate equation method gives the populations of the different species on the dust surface, which allows us to derive a simplified analytic expression for formation efficiencies of OH and H₂O that is easy to use in numerical codes. This means that the smallest set of reactions needed to reproduce the derived efficiencies from the rate equation calculation is determined. Although the analytic expression of the efficiencies is derived in Appendix A, the dependencies are given here:

$$\epsilon_{\text{OH}} = \frac{f_{\text{desOH}} n(\text{O}) n(\text{H}) \alpha_{\text{H}}}{R_{\text{acc}}(\text{O})} + \frac{f_{\text{desOH}'} n(\text{O}_3) n(\text{H}) \alpha_{\text{H}}}{R_{\text{acc}}(\text{O})} \quad (9)$$

$$\epsilon_{\text{H}_2\text{O}} = \frac{f_{\text{desH}_2\text{O}} n(\text{OH}) n(\text{H}) \alpha_{\text{H}}}{R_{\text{acc}}(\text{O})}. \quad (10)$$

The fractions f_{desOH} , $f_{\text{desOH}'}$, and $f_{\text{desH}_2\text{O}}$ are 0.36, 0.25, and 0.15, respectively. These are the fractions of the molecules that are released in the gas phase upon formation due to the exothermicity of the reaction. The accretion rate of oxygen is $R_{\text{acc}}(\text{O})$, while $n(\text{O})$, $n(\text{O}_3)$, and $n(\text{OH})$ are the dust surface densities of O, O₃, and OH, respectively, in monolayers, e.g., 1 monolayer = 100% coverage. These surface densities are listed in Appendix A. The variable α_{H} is the mobility of hydrogen atoms on the grain surface. The efficiency for the formation of water does *not* contain a term with accretion of OH from the gas phase, since the accretion of O followed by OH formation dominates the OH accretion.

A very important outcome of the simulation is that the formation of OH on grains is dominated by $\text{O} + \text{H} \rightarrow \text{OH}$ up to temperatures $T \sim 30\text{--}40 \text{ K}$ depending on the environment, and by $\text{O}_3 + \text{H} \rightarrow \text{OH} + \text{O}_2$ at higher temperatures. This implies that, besides expressions for $n(\text{O})$ and $n(\text{OH})$, those for $n(\text{O}_2)$ and $n(\text{O}_3)$ are also needed. For completeness, we also give the one for $n(\text{H}_2\text{O})$. The importance of the $\text{O}_3 + \text{H}$ reactions is shown in Fig. 6. It is obvious from the figure that the OH formation efficiency would be underestimated without considering the reaction with O₃. These results are consistent with those obtained by Ioppolo et al. (2008), where the formation of H₂O is dominated

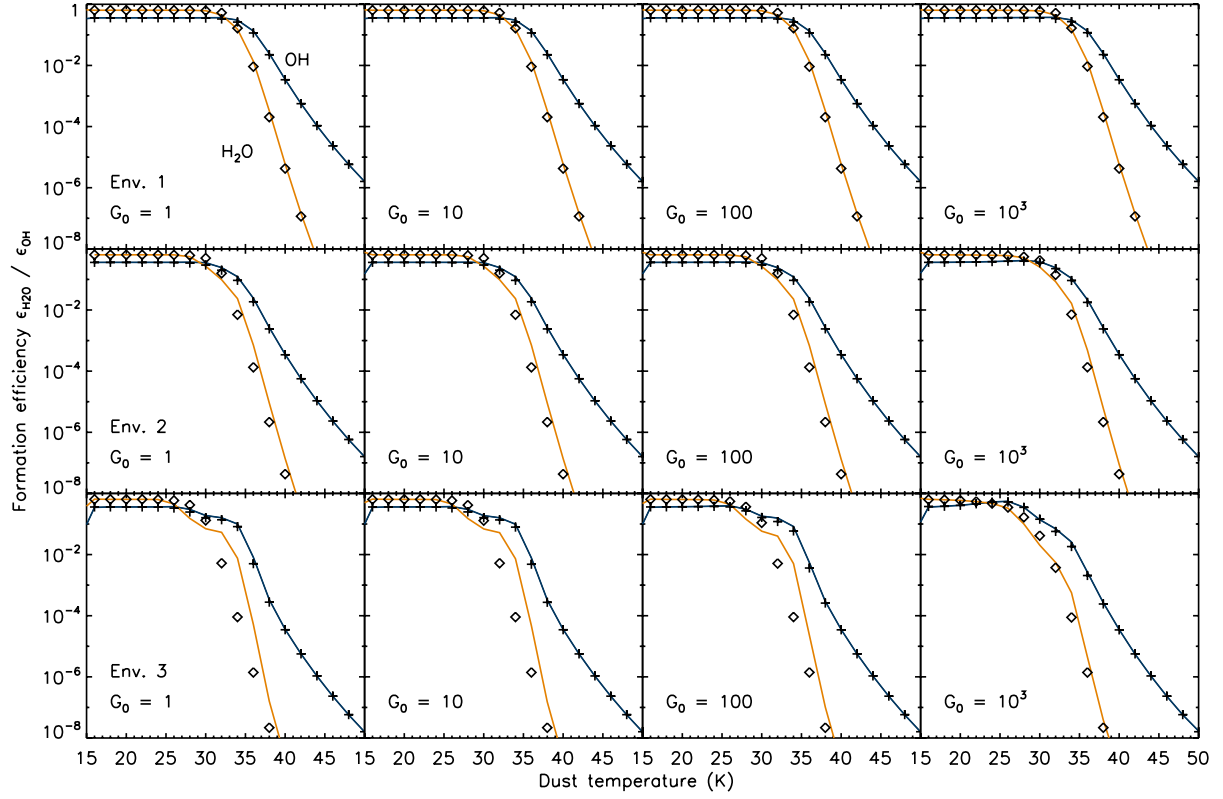


Fig. 2. Efficiencies for the formation of OH and H₂O in low density gas ($n = 10^3 \text{ cm}^{-3}$), in atomic and molecular hydrogen dominated gas. UV fields ranging from $G_0 = 1$ to 1000 are considered. Results from rate equations are shown by a solid and a dotted line for OH and H₂O, respectively. The analytical fit is shown by the crosses (OH) and diamonds (H₂O).

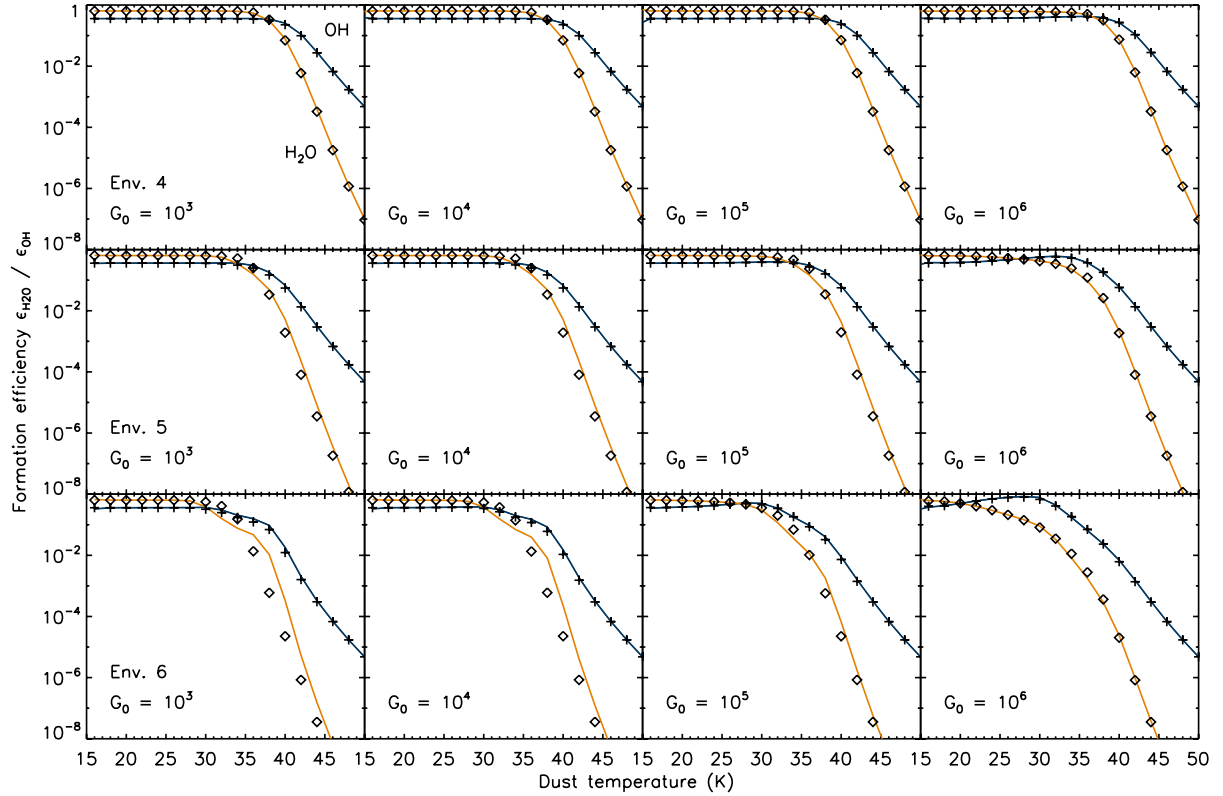


Fig. 3. Efficiencies for the formation of OH and H₂O in high-density gas ($n = 10^{5.5} \text{ cm}^{-3}$), in atomic and molecular dominated gas. UV fields ranging from $G_0 = 10^3$ to 10^6 are considered. Results from rate equations are shown by a solid line for OH and a dotted line for H₂O. The analytical fit is shown by the crosses (OH) and diamonds (H₂O).

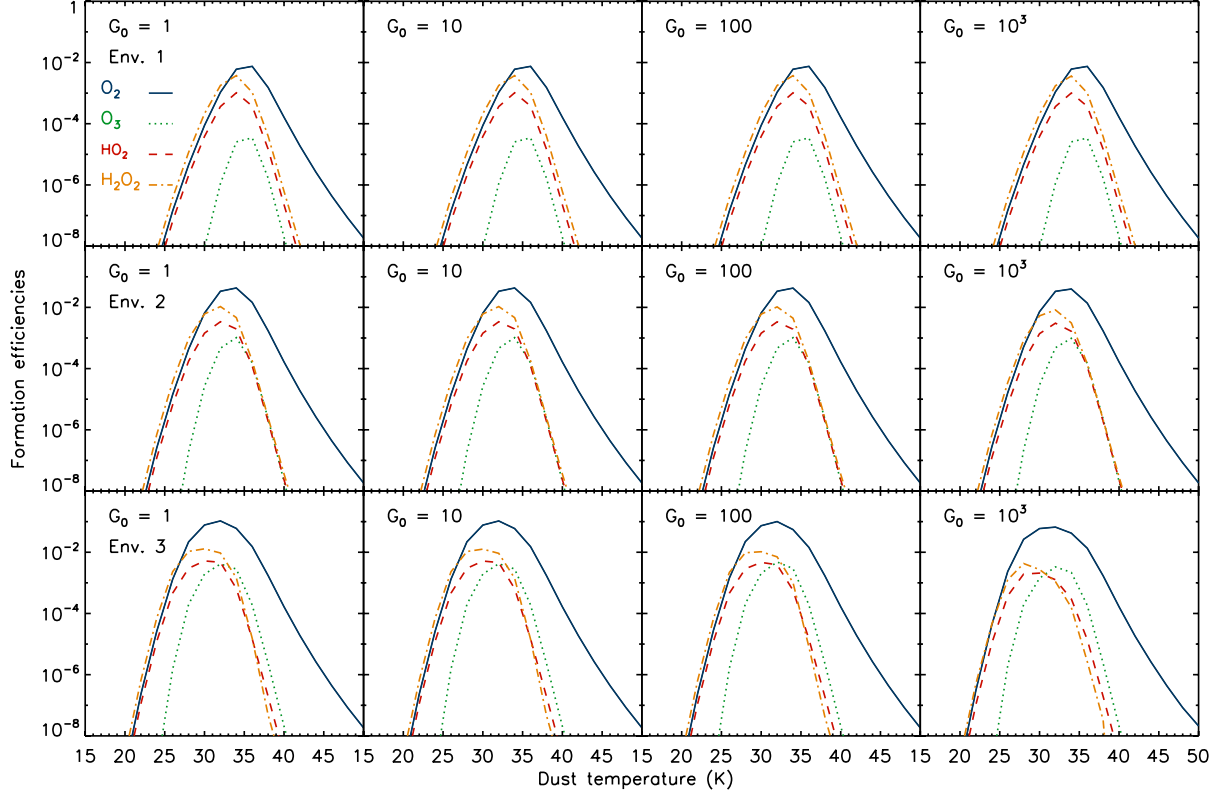


Fig. 4. Efficiencies for the formation of oxygen bearing species O₂, O₃, OH₂, and H₂O₂ in low-density gas ($n = 10^3 \text{ cm}^{-3}$), in atomic and molecular hydrogen dominated gas. UV fields ranging from $G_0 = 1$ to 1000 are considered.

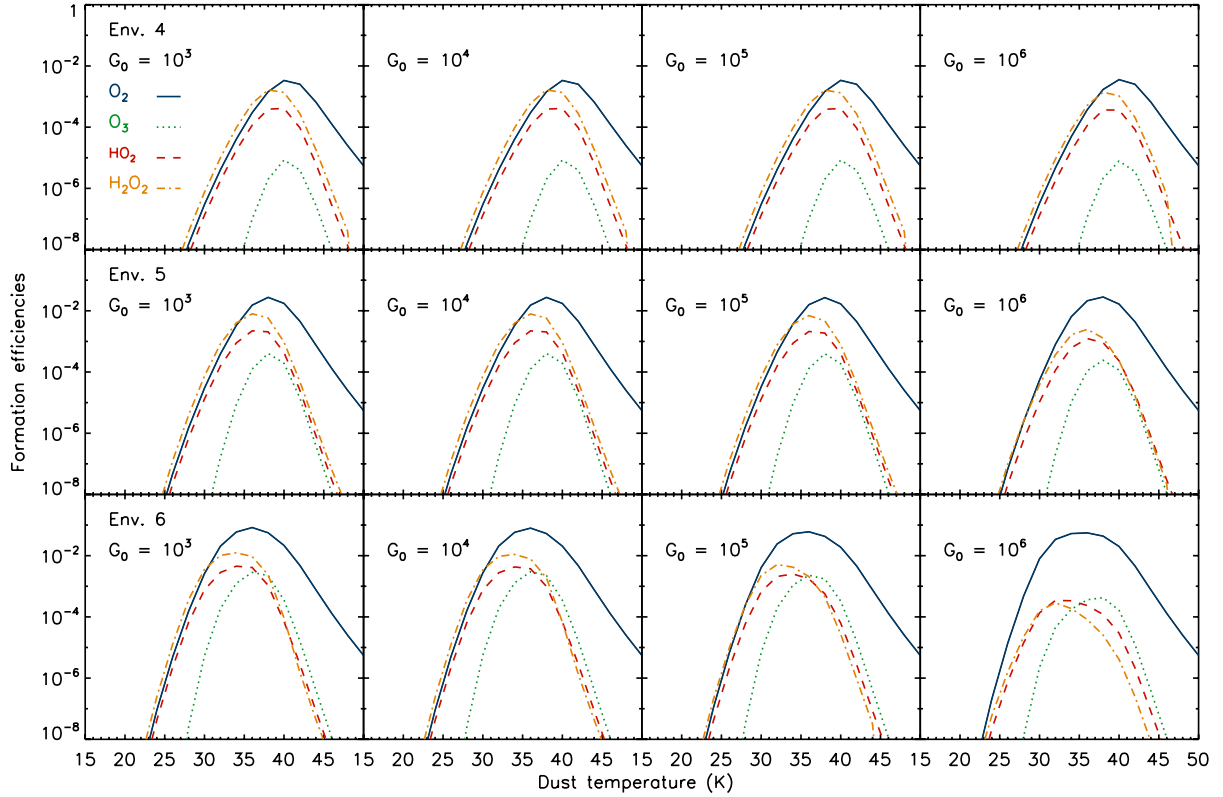


Fig. 5. Efficiencies for the formation of oxygen bearing species O₂, O₃, OH₂, and H₂O₂ in high-density gas ($n = 10^{5.5} \text{ cm}^{-3}$), in atomic and molecular dominated gas. UV fields ranging from $G_0 = 10^3$ to 10^6 are considered.

by the reaction sequence: $\text{H} + \text{O}_2 \rightarrow \text{HO}_2$, $\text{H} + \text{HO}_2 \rightarrow \text{H}_2\text{O}_2$ reaction rate paths in our work, but in this study, molecules are followed by $\text{H} + \text{H}_2\text{O}_2 \rightarrow \text{H}_2\text{O} + \text{OH}$. We do consider the same formed on *bare grains*, while Ioppolo et al. (2008) consider

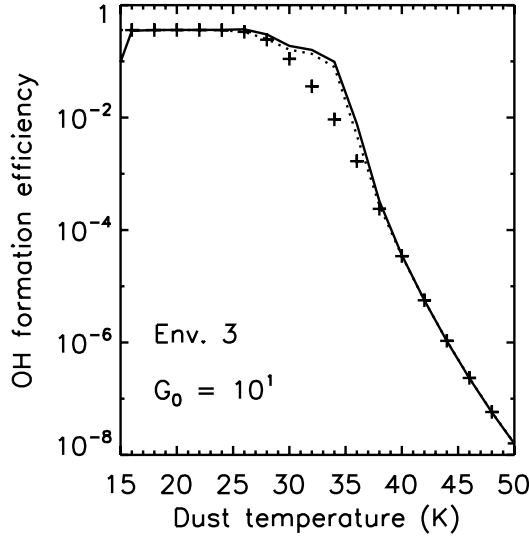


Fig. 6. Efficiencies for environment 3, largely molecular, low density $n = 10^3 \text{ cm}^{-3}$ gas, $G_0 = 10^4$: result from rate equation simulation (solid line), with (dotted line) and without (crosses) including of the reaction $\text{O}_3 + \text{H} \rightarrow \text{OH} + \text{O}_2$ in the analytical expression.

icy grain mantles. In our simulation a very significant part of the molecules are desorbed upon formation, while the cycle of photodissociation of H_2O and reformation on the grains significantly enhances the amount of water released in the gas phase. It does not require the complex route through H_2O_2 . There is also another way to form H_2O that involves O_3 . The reaction $\text{O} + \text{O}_2 \rightarrow \text{O}_3$ releases only a very small part into the gas phase upon formation (only a few percent), so although the efficiency for the formation of O_3 is low, there is a significant amount of O_3 on the surface to react to OH. The work by Hollenbach et al. (2009) is also very different from ours. They also treat a surface chemistry network, with H_2 , OH, O_2 , and some other different species, such as CH, CH_2 , and CH_3 . They do not include processes with O_3 and H_2O_2 , intermediate species in the processes leading to water formation. However, the biggest difference between our models are the dominant processes for delivering species back into the gas phase which are thermal, photo, and cosmic-ray desorption in the Hollenbach et al. (2009) paper, since they consider ices in that work. In our work, this is desorption upon formation. They consider a PDR environment, and therefore they allow for the formation of strongly bound icelayers. We assume that this does not occur because of the turbulent motions stripping the grains and the strong photodissociating radiation fields that are internally created by X-rays.

The approximate analytical fits to the rate equation results are very good for environments 1, 2, 4, and 5. However, environment 3 and 6, which are highly molecular, start to show deviations for dust temperatures $T_{\text{dust}} > 30 \text{ K}$ for the H_2O . This is due to additional reactions on the grain surface, that we did include in the rate equation method, but not in the approximate fit.

The fits should not be applied at very low dust temperatures ($T_{\text{dust}} < 10 \text{ K}$), because the analytical fits do not match the MC and rate equation results.

5. OH and H_2O formation: gas versus dust

To estimate the importance of OH and H_2O formation on dust grains, a comparison is made to the gas phase neutral-neutral

formation rates, which are given by (UMIST database: Le Teuff et al. 2000; Woodall et al. 2007)

$$k_{\text{OH,gas}} = 3.14 \times 10^{-13} \left(\frac{T_{\text{gas}}}{300} \right)^{2.70} \exp \left(-\frac{3150}{T_{\text{gas}}} \right) \text{ cm}^3 \text{ s}^{-1} \quad (11)$$

$$k_{\text{H}_2\text{O,gas}} = 2.05 \times 10^{-12} \left(\frac{T_{\text{gas}}}{300} \right)^{1.52} \exp \left(-\frac{1736}{T_{\text{gas}}} \right) \text{ cm}^3 \text{ s}^{-1}, \quad (12)$$

while the conservative dust phase reaction rates (using the cross sections from the MRN distribution, $n_{\text{dust}}\sigma/n(\text{H}) = 10^{-21} \text{ cm}^2$), are given by

$$k_{\text{OH,dust}} = 3.9 \times 10^{-17} \sqrt{\frac{T_{\text{gas}}}{100}} \epsilon_{\text{OH}}(T_{\text{dust}}) \text{ cm}^3 \text{ s}^{-1} \quad (13)$$

$$k_{\text{H}_2\text{O,dust}} = 3.9 \times 10^{-17} \sqrt{\frac{T_{\text{gas}}}{100}} \epsilon_{\text{H}_2\text{O}}(T_{\text{dust}}) \text{ cm}^3 \text{ s}^{-1}. \quad (14)$$

Depending on the environment (i.e., the ambient atomic and molecular abundances, gas and dust temperature, etc.), OH and H_2O form primarily on the dust grains or in the gas phase:

- Although the efficiency for formation of OH and H_2O on dust strongly depends on the dust temperature T_{dust} , the variations with gas temperature T_{gas} at a fixed dust temperature T_{dust} are very modest.
- The gas phase formation rates have high activation barriers of $T_{\text{gas}} = 3150$ and 1736 K for the formation of OH and H_2O , respectively. These channels are essentially cut off below temperatures $T_{\text{gas}} \lesssim 200 \text{ K}$.
- Formation on dust depends on abundances of O and H, and in the gas phase on H_2 , O, and OH.

For illustration purposes, a comparison between the gas and dust phase formation rates are shown for two different regimes in Fig. 7, a mostly atomic (left) and molecular (right) environment. We do not show the ion-molecule reaction rates, even though they are considered in our chemical network. These processes can be more efficient than the neutral reaction chain at low gas temperatures, which is a domain that we do not consider in this study. The abundances resemble the unshielded and shielded regions of model 4 in Meijerink & Spaans (2005) with density $n = 10^{5.5} \text{ cm}^{-3}$: the atomic environment has abundances $x_{\text{H}} = 1$, $x_{\text{H}_2} = 3 \times 10^{-4}$, $x_{\text{O}} = 3 \times 10^{-4}$, $x_{\text{OH}} = 10^{-8}$, 10^{-7} , and 10^{-6} (left), and the molecular environment has abundances $x_{\text{H}} = 10^{-2}$, $x_{\text{H}_2} = 0.5$, $x_{\text{O}} = 2 \times 10^{-4}$, and $x_{\text{OH}} = 10^{-6}$, 10^{-5} , and 10^{-4} (right). We adopt a range of OH abundances, to show the effect on the gas phase rates. We should note here that the OSU astrochemistry database¹ (maintained by Eric Herbst) claims a lower rate at high temperatures $T > 500$ for the $\text{H}_2 + \text{OH}$ reaction, and is given by $k_{\text{H}_2\text{O,gas}} = 8.4 \times 10^{-13} \exp(-1040/T_{\text{gas}})$. This would make the dust surface reaction even more important.

The efficiencies ϵ_{OH} and $\epsilon_{\text{H}_2\text{O}}$ are reasonably high, ~ 0.3 and ~ 0.6 , respectively, at dust temperatures $T_{\text{dust}} < 40 \text{ K}$. As a result, the H_2O dust formation rate is always dominating the gas phase formation in the atomic environment, while the OH dust formation is more important at temperatures $T_{\text{gas}} < 800 \text{ K}$. Even though OH and H_2O formation efficiencies drop very fast for higher dust temperatures, molecule formation with dust as catalyst will always become dominant below a certain gas temperature. The break-even point ranges from $T_{\text{gas}} \sim 100\text{--}400 \text{ K}$ depending on the dust temperature $T_{\text{dust}} = 45\text{--}90 \text{ K}$. We should

¹ <http://www.physics.ohio-state.edu/~eric/research.html>

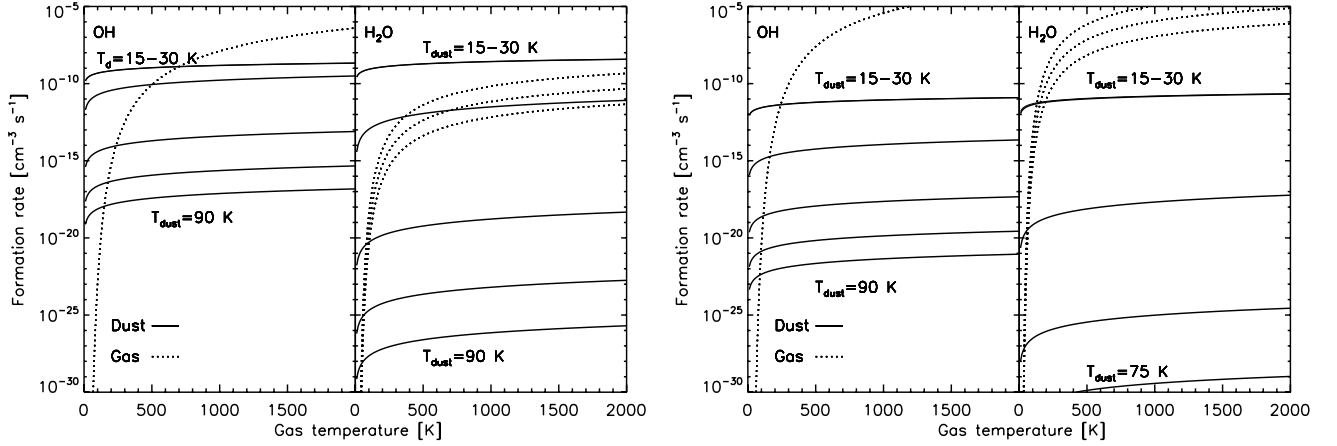


Fig. 7. Comparison of dust and gas phase reaction rates at $n = 10^{5.5} \text{ cm}^{-3}$ for an environment with low (left, $x_{\text{H}} = 1$, $x_{\text{H}_2} = 3 \times 10^{-4}$) and high molecular abundances (right, $x_{\text{H}} = 10^{-2}$, $x_{\text{H}_2} = 0.5$) as a function of gas temperature for dust temperatures $T_{\text{dust}} = 15, 30, 45, 60, 75$, and 90 K . The OH abundances are varied between $x_{\text{OH}} = 10^{-8}, 10^{-7}$, and 10^{-6} at low and $10^{-6}, 10^{-5}$, and 10^{-4} at high molecular abundances. This resembles the unshielded and shielded environment of model 4 in Meijerink & Spaans (2005).

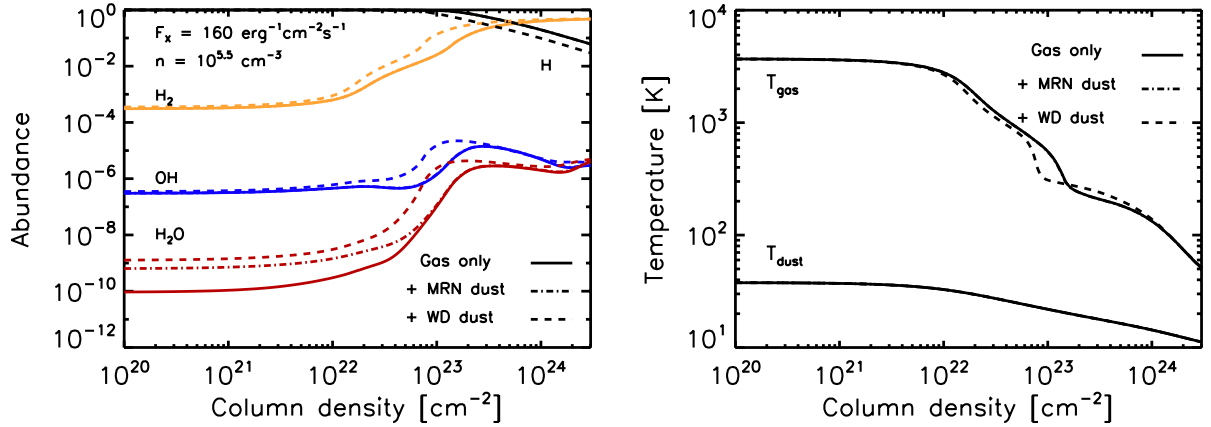


Fig. 8. X-ray irradiated cloud with density $n = 10^{5.5} \text{ cm}^{-3}$ and radiation field $F_X = 160 \text{ erg cm}^{-2} \text{ s}^{-1}$. (left) The abundances are shown for H, H₂, OH, and H₂O in three different cases: formation of OH and H₂O (i) in gas phase only (ii) in gas phase and on dust with an MRN size distribution, and (iii) in gas phase and on dust with a Weingartner & Draine size distribution. (right) Gas and dust temperatures for the same models. The temperature profiles for the gas only models and those where OH and H₂O formation on dust with an MRN distribution are included are the same.

keep in mind though that for $T_{\text{dust}} > 35\text{--}40 \text{ K}$, the formation rates are overestimated in our model.

This picture is different in a molecular environment. Even when the OH and H₂O formation efficiencies are high ($T_{\text{dust}} < 40 \text{ K}$), gas phase formation is more efficient at gas temperatures $T_{\text{gas}} > 100\text{--}200 \text{ K}$. Where in the atomic environment, the OH and H₂O gas phase formation is suppressed by the very low abundances of H₂, the dust phase formation rate is suppressed due to the low accretion rate of atomic hydrogen, which is depleted by two orders of magnitude. In the molecular environment, the formation of water in either the dust or gas phase is inefficient through neutral-neutral reactions, and ion-molecule reactions will also contribute significantly or even dominate, through the chain $\text{O}^+ + \text{H}_2 \rightarrow \text{OH}^+ + \text{H}$, $\text{OH}^+ + \text{H}_2 \rightarrow \text{H}_2\text{O}^+ + \text{H}$, $\text{H}_2\text{O}^+ + \text{H}_2 \rightarrow \text{H}_3\text{O}^+ + \text{H}$, followed by recombination $\text{H}_3\text{O}^+ + \text{e}^- \rightarrow \text{H}_2\text{O} + \text{H}$.

6. Application to X-ray environments

The inner regions of active galaxies from, e.g., Mrk 231, which are highly exposed to X-rays, are likely to exhibit enhanced

formation of H₂O on dust grains: (i) dust temperatures are expected to be moderate, $T_{\text{dust}} \sim 20\text{--}40 \text{ K}$, as the heating of dust is less efficient than by UV; (ii) the dynamical timescales are not very long and shocks are expected to be present, so dust particles are not expected to be or stay covered with ice; (iii) shocks may even be able to break up grains into smaller particles, therefore enlarging the effective surface of dust, where water is expected to form.

In Fig. 8, the chemical abundances of H, H₂, OH and H₂O and temperatures of an XDR model are shown with density $n = 10^{5.5} \text{ cm}^{-3}$, $F_X = 160 \text{ erg cm}^{-2} \text{ s}^{-1}$, and solar metallicity, $Z = Z_{\odot}$. Three different cases for the formation of OH and H₂O are considered:

1. Gas phase only: The OH abundance is $x_{\text{OH}} \sim 3 \times 10^{-7}$ at low column densities, rising to $\sim 10^{-5}$ just before the H/H₂ transition, leveling off to 10^{-6} . The H₂O abundance is $x_{\text{H}_2\text{O}} \sim 10^{-10}$ at column densities $N < 10^{22} \text{ cm}^{-2}$. As OH, H₂O becomes more and more abundant toward the H/H₂ transition, and reaches an abundance $x_{\text{H}_2\text{O}} \sim 10^{-6}$ beyond $N_{\text{H}} \sim 10^{23} \text{ cm}^{-2}$.

2. Gas and dust with an MRN distribution: nothing changes for OH. As shown in Fig. 7, the OH gas phase formation is more important at gas temperatures $T > 800$ K, at all dust temperatures. The abundance of H_2O , however, is higher by a factor 8, at low column densities $N_{\text{H}} < 10^{22} \text{ cm}^{-2}$, and thus the integrated column of warm water ($T \gtrsim 200$ K). At low temperatures (i.e., at high column densities, $N_{\text{H}} \gtrsim 10^{23} \text{ cm}^{-2}$), nothing changes for OH and H_2O , because the formation route is dominated by the ion-molecule reactions.
3. Gas and dust with a Weingartner & Draine distribution: the H_2 formation rate is now higher as well due to the higher surface area of the dust grains. As a result a small increase is seen in the H_2 abundance at low column densities. At the same time, the H/H_2 transition occurs at lower column densities, $N_{\text{H}} \sim 2 \times 10^{23} \text{ cm}^{-2}$, compared to $N_{\text{H}} \sim 3 \times 10^{23} \text{ cm}^{-2}$ compared to the gas phase model. As a result, the OH also peaks at a lower column density, and reaches a slightly higher peak abundance. *The OH enhancement is due to a higher gas phase formation rate, as a result of the higher H_2 abundance, not because of the higher formation rate on dust.* The water abundance in the high temperature range is enhanced by more than an order of magnitude. As the H/H_2 transition occurs at lower column densities, the rise in the water abundance also occurs at these lower column densities.

To summarize, we find that the warm H_2O column density is enhanced by approximately an order of magnitude over a column $N_{\text{H}} \sim 5 \times 10^{22} \text{ cm}^{-2}$, when including H_2O formation on dust. At higher column densities, where temperatures decrease and the gas becomes molecular, the formation on dust grains does not affect the abundance of H_2O . The OH abundance is only indirectly affected by the higher H_2 formation rate on dust when using the Weingartner & Draine (2001) distribution with the higher dust cross section.

7. Discussion and conclusions

The role and importance of OH and H_2O formation on dust was investigated. We used both Monte Carlo and rate equation simulations to determine the efficiency at which oxygen is converted into OH and H_2O on dust grains. The rate equation method is very fast, and is used to calculate a grid of different environments, both low and high density ones ($n = 10^3$ and $10^{5.5} \text{ cm}^{-3}$), and different molecular fractions, radiation fields, and temperatures (see Figs. 2 and 3). The Monte Carlo simulations are used as a consistency check, and because these simulations are fairly slow, efficiencies are only calculated for a limited number of dust temperatures. At dust temperatures $T_{\text{dust}} < 35\text{--}40$ K, the Monte Carlo and rate equation method yield the same results, within the uncertainties. At higher dust temperatures, the efficiencies obtained by the two methods start to deviate: the rate equation method systematically obtains higher efficiencies than the MC method. This is where the process of OH and H_2O formation on dust enters the stochastic regime. We fitted analytical expression to the derived results from the rate equations, such that they are easy to implement in any chemistry code. We find the following:

- OH formation on grains and release to gas phase is dominated by $\text{O} + \text{H} \rightarrow \text{OH}$ ($T_{\text{dust}} < 35\text{--}40$ K and by $\text{O}_3 + \text{H} \rightarrow \text{OH} + \text{O}_2$ (for higher dust temperatures). The high dust temperatures result in decreasing hydrogen surface densities, because of the increasing evaporation rates.

- H_2O formation on grains and release to gas phase is dominated by $\text{OH} + \text{H} \rightarrow \text{H}_2\text{O}$.
- The fraction of OH that is released in the gas upon formation is on the order of 36% and for H_2O it is 9%. However, in strong UV radiation fields, water on the dust can be photodissociated, and then reform to water, which then gives another contribution to the gas phase. This loop of forming-dissociating-forming increases the effective fraction of H_2O released in the gas by a factor 5.
- A comparison of gas and dust phase formation rates (see Fig. 7) shows that formation on dust grains is important for dust temperatures lower than 40 K.

We implemented the analytical expressions for the formation rates in the Meijerink & Spaans (2005) XDR code, and calculated a model with parameters representative of the inner regions of active galaxies. We find that the additional processes enhances the integrated column density of warm water. This increased column helps explain unusually strong water lines observed for Mrk 231, which has an accreting supermassive black hole with X-ray luminosity $L_{\text{X}} \sim 10^{44} \text{ erg s}^{-1}$. This ULIRG shows, for example, a very high $\text{H}_2\text{O } 2_{02}\text{--}1_{11}/\text{CO } J = 8\text{--}7$ ratio. These ratios are not observed in typical starforming environments that are mainly exposed to UV-only (PDR) environments, such as the one observed in the Orion Bar (Habart et al. 2010) or M 82 (Panuzzo et al. 2010), where this ratio is typically an order of magnitude or more less than in Mrk 231, and as such these very bright H_2O lines might be typical of regions that are exposed to high X-ray fluxes, close to the accreting black hole in active galaxies.

Acknowledgements. We thank the anonymous referee for a careful reading of the manuscript and a constructive report that significantly improved the paper.

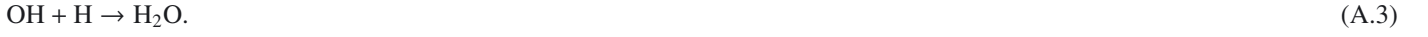
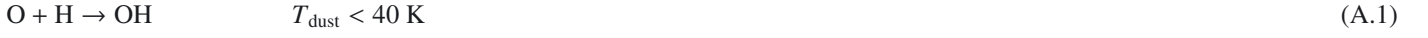
References

- Akai, Y., & Saito, S. 2003, Jpn. J. Appl. Phys., 42, 640
 Andersson, S., & van Dishoeck, E. F. 2008, A&A, 491, 907
 Andersson, S., Al-Halabi, A., Kroes, G.-J., & van Dishoeck, E. F. 2006, J. Chem. Phys., 124, 064715
 Bergeron, H., Rougeau, N., Sidis, V., Teillet-Billy, D., & Aguilon, F. 2008, J. Phys. Chem. A, 112, 11921
 Bokkenberg, A., Carswell, R. F., Allen, D. A., et al. 1977, MNRAS, 178, 451
 Braitto, V., Della Ceca, R., Piconcelli, E., et al. 2004, A&A, 420, 79
 Cazaux, S., & Tielens, A. G. G. M. 2002, ApJ, 575, L29
 Cazaux, S., & Tielens, A. G. G. M. 2004, ApJ, 604, 222
 Cazaux, S., Cobut, V., Marseille, M., Spaans, M., & Caselli, P. 2010, A&A, 522, A74
 Cuppen, H. M., & Herbst, E. 2007, ApJ, 668, 294
 Cuppen, H. M., Ioppolo, S., Romanzin, C., & Linnartz, H. 2010a, Phys. Chem. Chem. Phys., 12, 12077
 Cuppen, H. M., Kristensen, L. E., & Gavardi, E. 2010b, MNRAS, 406, L11
 Draine, B. T. 1995, Ap&SS, 233, 111
 Fischer, J., Sturm, E., González-Alfonso, E., et al. 2010, A&A, 518, L41
 Galli, D., & Palla, F. 1998, A&A, 335, 403
 Glover, S. C. O. 2003, ApJ, 584, 331
 González-Alfonso, E., Fischer, J., Isaak, K., et al. 2010, A&A, 518, L43
 Habart, E., Dartois, E., Abergel, A., et al. 2010, A&A, 518, L116
 Hollenbach, D. J., & Tielens, A. G. G. M. 1999, Rev. Mod. Phys., 71, 173
 Hollenbach, D., Kaufman, M. J., Bergin, E. A., & Melnick, G. J. 2009, ApJ, 690, 1497
 Hornekar, L., Rauls, E., Xu, W., et al. 2006, Phys. Rev. Lett., 97, 186102
 Ioppolo, S., Cuppen, H. M., Romanzin, C., van Dishoeck, E. F., & Linnartz, H. 2008, ApJ, 686, 1474
 Ioppolo, S., Cuppen, H. M., Romanzin, C., van Dishoeck, E. F., & Linnartz, H. 2010, Phys. Chem. Chem. Phys., 12, 12065

- Jeleva, A., Marinelli, F., Ferro, Y., Allouche, A., & Brosset, C. 2008, *J. Phys. Chem. A*, 112, 11921
- Katz, N., Furman, I., Biham, O., Pirronello, V., & Vidali, G. 1999, *ApJ*, 522, 305
- Launay, J. M., Le Dourneuf, M., & Zeppen, C. J. 1991, *A&A*, 252, 842
- Le Teuff, Y. H., Millar, T. J., & Markwick, A. J. 2000, *A&AS*, 146, 157
- Lee, G., Lee, B., Kim, J., & Cho, K. 2009, *J. Phys. Chem. C*, 113, 14225
- Maloney, P. R., Hollenbach, D. J., & Tielens, A. G. G. M. 1996, *ApJ*, 466, 561
- Mathis, J. S., Rimpl, W., & Nordsieck, K. H. 1977, *ApJ*, 217, 425
- Meijerink, R., & Spaans, M. 2005, *A&A*, 436, 397
- Miyauchi, N., Hidaka, H., Chigai, T., et al. 2008, *Chem. Phys. Lett.*, 456, 27
- Mokrane, H., Chaabouni, H., Accolla, M., et al. 2009, *ApJ*, 705, L195
- Morisset, S., Aguilon, F., Sizun, M., & Sidis, V. 2004, *J. Chem. Phys.*, 121, 6493
- Neufeld, D. A., & Kaufman, M. J. 1993, *ApJ*, 418, 263
- Neufeld, D. A., Lepp, S., & Melnick, G. J. 1995, *ApJS*, 100, 132
- Nisini, B., Benedettini, M., Giannini, T., et al. 1999, *A&A*, 350, 529
- Panuzzo, P., Rangwala, N., Rykala, A., et al. 2010, *A&A*, 518, L37
- Papadopoulos, P. P., van der Werf, P., Isaak, K., & Xilouris, E. M. 2010, *ApJ*, 715, 775
- Pirronello, V., Liu, C., Roser, J. E., & Vidali, G. 1999, *A&A*, 344, 681
- Sha, X. 2002, *Surface Science*, 496, 318
- Sha, X., Jackson, B., Lemoine, D., & Lepetit, B. 2005, *J. Chem. Phys.*, 122, 014709
- Sturm, B., Bouwman, J., Henning, T., et al. 2010, *A&A*, 518, L129
- van der Werf, P. P., Isaak, K. G., Meijerink, R., et al. 2010, *A&A*, 518, L42
- van Kempen, T. A., Green, J. D., Evans, N. J., et al. 2010, *A&A*, 518, L128
- Wagner, A. F., & Graff, M. M. 1987, *ApJ*, 317, 423
- Weingartner, J. C., & Draine, B. T. 2001, *ApJ*, 548, 296
- Woodall, J., Agúndez, M., Markwick-Kemper, A. J., & Millar, T. J. 2007, *A&A*, 466, 1197
- Zecho, T., Guttler, A., Sha, X., Jackson, B., & Kuppers, J. 2002, *J. Chem. Phys.*, 117, 8486

Appendix A: OH and H₂O formation efficiencies

The formation efficiency with which gas-phase OH and H₂O are formed can be written as a balance between the desorption fraction from the grain upon formation and the accretion rate of oxygen atoms onto the grain. On grains, two formation processes play a major role for OH in different dust temperature regimes, while for H₂O, there is only one major formation route:



Therefore, the efficiencies should be written as

$$\epsilon_{\text{OH}} = \frac{f_{\text{desOH}} n(\text{O}) n(\text{H}) \alpha_{\text{H}}}{R_{\text{acc}}(\text{O})} + \frac{f_{\text{desOH}} n(\text{O}_3) n(\text{H}) \alpha_{\text{H}}}{R_{\text{acc}}(\text{O})} \quad (\text{A.4})$$

$$\epsilon_{\text{H}_2\text{O}} = \frac{f_{\text{desH}_2\text{O}} n(\text{OH}) n(\text{H}) \alpha_{\text{H}}}{R_{\text{acc}}(\text{O})}, \quad (\text{A.5})$$

where $n(\text{H})$, $n(\text{O})$, $n(\text{O}_3)$, and $n(\text{OH})$ are the surface density of H, O, O₃, and OH on the dust. These hydrogen and oxygen atoms are accreted from the gas phase at a rate of

$$R_{\text{acc}}(\text{H}) = \frac{n_{\text{HI}} v_{\text{HI}}}{N_{\text{S}}} \text{ monolayers s}^{-1}, \quad (\text{A.6})$$

$$R_{\text{acc}}(\text{O}) = \frac{n_{\text{OI}} v_{\text{OI}}}{N_{\text{S}}} \text{ monolayers s}^{-1}, \quad (\text{A.7})$$

where n_{HI} and n_{OI} are the densities of atomic hydrogen and oxygen in the gas phase, v_{HI} and v_{OI} the mean velocities of these atoms:

$$v_{\text{HI}} \sim 1.45 \times 10^5 \sqrt{\frac{T_{\text{gas}}}{100}} \text{ cm s}^{-1}, \quad (\text{A.8})$$

$$v_{\text{OI}} \sim 3.6 \times 10^4 \sqrt{\frac{T_{\text{gas}}}{100}} \text{ cm s}^{-1}, \quad (\text{A.9})$$

and $N_{\text{S}} = \frac{1}{a_{\text{pp}}^2} = 8 \times 10^{14} \text{ cm}^{-2}$ the number of sites per cm^2 . The H atoms arrive on the dust surface in physisorbed sites, and then continue to a chemisorbed site or evaporate from the surface. As a result, the density on the surface can be written as

$$n(\text{H}) = \frac{R_{\text{acc}}(\text{H})}{\alpha_{\text{pc}} + R_{\text{evapH}}}. \quad (\text{A.10})$$

The mobility for hydrogen atoms to enter into a chemisorbed site is mostly due to tunneling, which we can then approximate by $\alpha_{\text{pc}} = 10^{12} \exp(-2a_{\text{pc}} \sqrt{\frac{2m_{\text{H}} E_{\text{A}}}{\hbar^2}}) = 6.7 \times 10^{-3}$, where a_{pc} and E_{A} are the width and the height of the barrier between physisorbed and chemisorbed sites. The adopted values are $a_{\text{pc}} = 1.5 \text{ \AA}$ and $E_{\text{A}} = 2870 \text{ K}$. The hydrogen surface density can then be calculated as

$$n(\text{H}) = \frac{1.45 \times 10^{-10} n_{\text{HI}} \sqrt{\frac{T_{\text{gas}}}{100}}}{6.7 \times 10^{-3} + 10^{12} \exp(\frac{-526}{T_{\text{dust}}})}. \quad (\text{A.11})$$

The formation of OH and H₂O can be determined analytically from the rate equations, when the most important processes are considered. The rate equations of O, OH, and H₂O, as well as O₂ and O₃, are considered here, since the efficiency of OH depends on the abundance of O₃ as well. The abundance change as a function of time for these species are given by the following expressions:

$$\frac{dn(\text{O})}{dt} = R_{\text{acc}}(\text{O}) - n(\text{O}) n(\text{H}) \alpha_{\text{H}} + n(\text{OH}) R_{\text{photOH}} - 2n(\text{O})^2 \alpha_{\text{O}} - R_{\text{evapO}} n(\text{O}) \quad (\text{A.12})$$

$$\frac{dn(\text{OH})}{dt} = (1 - f_{\text{desOH}}) n(\text{O}) n(\text{H}) \alpha_{\text{H}} + R_{\text{photH}_2\text{O}} n(\text{H}_2\text{O}) - R_{\text{photOH}} n(\text{OH}) - n(\text{OH}) n(\text{H}) \alpha_{\text{H}} - R_{\text{evapOH}} n(\text{OH}) \quad (\text{A.13})$$

$$\frac{dn(\text{H}_2\text{O})}{dt} = (1 - f_{\text{desH}_2\text{O}}) n(\text{OH}) n(\text{H}) \alpha_{\text{H}} - R_{\text{photH}_2\text{O}} n(\text{H}_2\text{O}) - n(\text{H}_2\text{O}) n(\text{H}) \alpha_{\text{H}} \quad (\text{A.14})$$

$$\frac{dn(\text{O}_2)}{dt} = (1 - f_{\text{desO}_2}) n(\text{O})^2 \alpha_{\text{O}} + (1 - f'_{\text{desO}_2}) n(\text{OH}) n(\text{O}) \alpha_{\text{OH}} - n(\text{O}_2) n(\text{H}) \alpha_{\text{H}} - R_{\text{photO}_2} n(\text{O}_2) - R_{\text{evapO}_2} n(\text{O}_2) \quad (\text{A.15})$$

$$\frac{dn(\text{O}_3)}{dt} = (1 - f_{\text{desO}_3}) n(\text{O}_2) n(\text{O}) \alpha_{\text{O}} - n(\text{O}_3) n(\text{H}) \alpha_{\text{H}} - n(\text{O}_3) n(\text{OH}) \alpha_{\text{OH}} - R_{\text{photO}_3} n(\text{O}_3). \quad (\text{A.16})$$

Table A.1. Adopted photodissociation rates in the approximation.

Reaction	Rate (s ⁻¹)
OH + photon → O + H	3.9(-10)
O ₂ + photon → O + O	7.9(-10)
H ₂ O + photon → OH + H	8.0(-10)
O ₃ + photon → O ₂ + O	1.9(-09)

When we consider the surface densities of the species at steady state, the surface densities of the species become

$$n(\text{O}) = \frac{R_{\text{acc}}(\text{O})}{n(\text{H})\alpha_{\text{H}} + R_{\text{evap}_\text{O}} - R_{\text{phot}_{\text{OH}}} \frac{1-f_{\text{des}_{\text{OH}}}}{f_{\text{des}_{\text{H}_2\text{O}}} + \frac{R_{\text{phot}_{\text{OH}}}}{n(\text{H})\alpha_{\text{H}}}} + \sqrt{2R_{\text{acc}}(\text{O})\alpha_{\text{O}}} \quad (\text{A.17})$$

$$n(\text{OH}) = \frac{(1 - f_{\text{des}_{\text{OH}}})n(\text{O})}{f_{\text{des}_{\text{H}_2\text{O}}} + \frac{R_{\text{phot}_{\text{OH}}}}{n(\text{H})\alpha_{\text{H}}} + \frac{R_{\text{evap}_{\text{OH}}}}{n(\text{H})\alpha_{\text{H}}}} \quad (\text{A.18})$$

$$n(\text{H}_2\text{O}) = \frac{(1 - f_{\text{des}_{\text{H}_2\text{O}}})n(\text{OH})n(\text{H})\alpha_{\text{H}}}{n(\text{H})\alpha_{\text{H}} + R_{\text{phot}_{\text{H}_2\text{O}}}} \quad (\text{A.19})$$

$$n(\text{O}_2) = \frac{(1 - f_{\text{des}_{\text{O}_2}})n(\text{O})^2\alpha_{\text{O}} + (1 - f'_{\text{des}_{\text{O}_2}})n(\text{OH})n(\text{O})\alpha_{\text{OH}}}{n(\text{H})\alpha_{\text{H}} + R_{\text{phot}_{\text{O}_2}} + R_{\text{evap}_{\text{O}_2}}} \quad (\text{A.20})$$

$$n(\text{O}_3) = \frac{(1 - f_{\text{des}_{\text{O}_3}})n(\text{O}_2)n(\text{O})\alpha_{\text{O}}}{n(\text{H})\alpha_{\text{H}} + n(\text{OH})\alpha_{\text{OH}} + R_{\text{phot}_{\text{O}_3}}} \quad (\text{A.21})$$

Filling in the equations for the various included reactions, the expressions become as follows:

$$n(\text{O}) = \frac{R_{\text{acc}}(\text{O})}{10^{12} \exp\left(\frac{-350}{T_{\text{dust}}}\right)n(\text{H}) + 10^{12} \exp\left(\frac{-1366}{T_{\text{dust}}}\right) - \frac{3.90 \times 10^{-10} G_0 \cdot 0.64}{0.15 + \frac{3.9 \times 10^{-10} G_0}{n(\text{H}) \exp\left(\frac{-350}{T_{\text{dust}}}\right)}} + \sqrt{R_{\text{acc}}(\text{O}) 10^{12} \exp\left(\frac{-910}{T_{\text{dust}}}\right)}} \quad (\text{A.22})$$

$$n(\text{OH}) = \frac{0.64n(\text{O})}{0.15 + \frac{3.90 \times 10^{-10} G_0}{n(\text{H}) 10^{12} \exp\left(\frac{-350}{T_{\text{dust}}}\right)} + \frac{10^{12} \exp\left(\frac{-1336}{T_{\text{dust}}}\right)}{n(\text{H}) 10^{12} \exp\left(\frac{-350}{T_{\text{dust}}}\right)}} \quad (\text{A.23})$$

$$n(\text{H}_2\text{O}) = \frac{0.85n(\text{H})n(\text{OH}) 10^{12} \exp\left(\frac{-350}{T_{\text{dust}}}\right)}{n(\text{H}) 10^{12} \exp\left(\frac{-350}{T_{\text{dust}}}\right) + 8 \times 10^{-10} G_0} \quad (\text{A.24})$$

$$n(\text{O}_2) = \frac{0.975n(\text{OH})n(\text{O}) 10^{12} \exp\left(\frac{-890}{T_{\text{dust}}}\right) + 0.64n(\text{O}_2) 10^{12} \exp\left(\frac{-910}{T_{\text{dust}}}\right)}{n(\text{H}) 10^{12} \exp\left(\frac{-350}{T_{\text{dust}}}\right) + 7.90 \times 10^{-10} G_0 + 10^{12} \exp\left(\frac{-1416}{T_{\text{dust}}}\right)} \quad (\text{A.25})$$

$$n(\text{O}_3) = \frac{0.977n(\text{O})n(\text{O}_2) 10^{12} \exp\left(\frac{-910}{T_{\text{dust}}}\right)}{n(\text{H}) 10^{12} \exp\left(\frac{-350}{T_{\text{dust}}}\right) + 10^{12} \exp\left(\frac{-890}{T_{\text{dust}}}\right)n(\text{OH}) + 1.8 \times 10^{-10} G_0} \quad (\text{A.26})$$